Is Nature fundamentally continuous or discrete, and what violations of the Null Energy Condition tell us about information exchange between prior to present universes ?

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Abstract

Our contention, is that reality is actually analog, but that at a critical limit, as when the Octonian gravity condition kicks in, that for a time it is made to appear discrete. This due to an initial phase transition just at the start of the big bang. Our second consideration is, that symmetry breaking models, i.e. the Higgs boson are in themselves not appropriate or necessary for the formation of particles with mass just before Octonionic gravity which could arise in pre Planckian physics models without a potential. Also the necessity of potentials for pre Octonionic gravity physics can be circumvented via judicious use of Sherrer k essence physics The universe is now dominated by DE leading to renewed acceleration, so what we can do is to examine how DE arose in the first place, and what role cosmologies not obeying the null energy condition play in terms of facilitating information exchange from a prior to the present universe.

Introduction

Our presentation takes note of several developments. First of all a feed into cosmological vacuum energy has been modeled, and that we have ideas as to how to inter relate four and five dimensional vacuum energies. Secondly, a mechanism for the onset of Octonian gravity is stated, as a consequence as to a build up of a peak temperature for its inception, at the time space time flattens. The onset of pre Octonionic gravity, with tiny masses associated with gravitons, is in line with Quantum mechanics as embedded within a larger, non linear classical theory (I.e. go to the Pilot Model, to get an idea of what is involved. That plus t'Hoofts deterministic quantum mechanics construction) Thirdly, we suggest that the transition from highly curved space time , which is pre Octonian gravity , ie. non quantum state, to quantum state, is due to a chaotic mapping which we present in this document. That chaotic mapping also has that there would be up to Planckian space time an explosion of the degrees of freedom. I.e. this degree of freedom explosion would be where we obtain quantum dynamics. Thermal inputs for the push to quantum dynamics are the first topic brought up for our perusal of this document.

Not only is there a chaotic mapping, and explosion as to the number of degrees of freedom, but the entire construction allows us to investigate the entire premise of the null energy condition, and what happens if it is violated in the beginning. Our investigations indicate, that if, as stated by Steinhardt , and Wesley [1] , one has cosmologies congruent with the null energy condition, with inflation only viable if the initial dark energy phase consistent with observations only possible if both Newton's gravitational constant and the dark energy equation-of-state vary with time. What we have is a dark energy time varying equation of state, due to vacuum energy being altered as below. IF we do not put in Newton's constant changing with time, we then have very strange cosmological behavior which may nix out the idea of not only initial dark energy but the entire working hypothesis of exchange of information from a prior to the present universe.

We refer to the concept of the null energy condition toward the end of the document and how we can then look at if it is kept, or violated as part of the confirmation of if conditions exist for the Null energy condition. Furthermore, we will also state how issues connected with the null energy condition are of essential importance as to if information can be exchanged between a prior to the present universe. We will state why , and make recommendations as to how confirmation of this last point can be tested via data.

Vacuum energy , sources and commentary

Begin first with looking at different value of the cosmological vacuum energy parameters, in four and five dimensions [2]

$$\left|\Lambda_{5-\dim}\right| \approx c_1 \cdot \left(1/T^{\alpha}\right) \qquad (1)$$

in contrast with the more traditional four-dimensional version of the same, minus the minus sign of the brane world theory version. The five-dimensional version is actually connected with Brane theory and higher dimensions, whereas the four-dimensional version is linked to more traditional De Sitter space-time geometry, as given by Park (2003) [3]

$$\Lambda_{4-\dim} \approx c_2 \cdot T^{\beta} \qquad (2)$$

If one looks at the range of allowed upper bounds of the cosmological constant, the difference between what Barvinsky (2006) [4] recently predicted, and Park (2003) [3] is:

$$\Lambda_{4-\dim} \propto c_2 \cdot T^{\beta} \xrightarrow[graviton-production-as-time>t(Planck)]{} 360 \cdot m_P^2 << c_2 \cdot \left[T \approx 10^{32} K\right]^{\beta} \qquad (3)$$

Right after the gravitons are released, one still sees a drop-off of temperature contributions to the cosmological constant .Then one can write, for small time values $t \approx \delta^1 \cdot t_P$, $0 < \delta^1 \leq 1$ and for temperatures sharply lower than $T \approx 10^{12}\, Kelvin$, Beckwith (2008), where for a positive integer n [5]

$$\frac{\Lambda_{4-\dim}}{\left|\Lambda_{5-\dim}\right|} - 1 \approx \frac{1}{n} \qquad (4)$$

If there is an order of magnitude equivalence between such representations, there is a quantum regime of gravity that is consistent with fluctuations in energy and growth of entropy. An order-of-magnitude estimate will be used to present what the value of the vacuum energy should be in the neighborhood of Planck time in the advent of nucleation of a new universe. The significance of Eq (4) is that at very high temperatures, it re enforces what the author brought up with Tigran Tchrakian, in Bremen,[6] August 29th, 2008. I.e., one would like to have a uniform value of the cosmological constant in the gravitating Yang-Mills fields in quantum gravity in order to keep the gauges associated with instantons from changing. When one has, especially for times $t_1, t_2 <$ Planck time t_P and $t_1 \neq t_2$, with temperature $T(t_1) \neq T(t_2)$, then $\Lambda_4(t_1) \neq \Lambda_4(t_2)$. I.e., in the regime of high temperatures, one has $T(t_1) \neq T(t_2)$ for times $t_1, t_2 <$ Planck time t_P and $t_1 \neq t_2$, such that gauge invariance necessary for soliton (instanton) stability is broken [6]. That breaking of instanton stability due to changes of $\Lambda_4(t_1) \neq \Lambda_4(t_2)$ will be our point of where we move from an embedding of quantum mechanics in an analog reality, to the quantum regime. I.e. as one reaches to high temperature, analog reality mimics digital quantum mechanics. Let us now look at different characterizations of the discontinuity, which is the boundary between analog reality, and Octonian gravity. First of all, one can look at scale factor evolution.

What leads to causal discontinuity in scale factor evolution?

The Friedmann equation [7] for the evolution of a scale factor $a(t)$,

$$\left(\dot{a}/a\right)^2 = \frac{8\pi G}{3} \cdot \left[\rho_{rel} + \rho_{matter}\right] + \frac{\Lambda}{3} \qquad (5)$$

suggests a non-partially ordered set evolution of the scale factor with evolving time, thereby implying a causal discontinuity. The validity of this formalism is established by rewriting the Friedman equation as follows: $a(t^*) < l_P$ for $t^* < t_P =$ Planck time, and $a_0 \equiv l_P$, for a discrete equation model of Eq (6) [5]

$$\left[\frac{a(t^* + \delta t)}{a(t^*)}\right] - 1 <$$

$$\frac{\left(\delta t \cdot l_P\right)}{\left(\sqrt{3/8\pi\Lambda}\right)} \cdot \left[\frac{1}{24\pi \cdot a^2(t^*)} + \frac{1}{\Lambda} \cdot \left[(\rho_{rel})_0 \cdot \frac{a_0^4}{a^6(t^*)} + (\rho_m)_0 \cdot \frac{a_0^3}{a^5(t^*)}\right]\right]^{1/2} \tag{6}$$

$$\xrightarrow[\delta t \to \varepsilon^+, \Lambda \neq \infty, a \neq 0]{} \left(\frac{\delta t \cdot [l_P / a(t^*)]}{\sqrt{3/8\pi}}\right) \cdot \sqrt{\frac{(\rho_{rel})_0 a_0^4}{a^4(t^*)} + \frac{(\rho_m)_0 a_0^3}{a^3(t^*)}} \approx \varepsilon^+ << 1$$

So in the initial phases of the big bang, with very large vacuum energy $\neq \infty$ and $a(t^*) \neq 0, 0 < a(t^*) << 1$, the following relation, which violates (signal) causality, is obtained for very small fluctuation $a(t^*) < l_P$ for $t^* < t_P =$ Planck time, and $a_0 \neq l_P, a_0 >> l_P$, which indicates that [7]

$$\rho_{rel} \equiv \left(\frac{a_{present-era}}{a(t)}\right)^4 \cdot (\rho_{rel})_{present-era} \tag{7}$$

And

$$\rho_m \equiv \left(\frac{a_{present-era}}{a(t)}\right)^3 \cdot (\rho_m)_{present-era} \tag{8}$$

Using the above equation creates the following as plausible estimates, which can be reviewed, as needed. For large, but not infinite temperatures, and for $\Lambda \sim c_1 T^{\alpha}$ [5]

$$\left(\frac{\delta t \cdot [l_P / a(t^*)]}{\sqrt{3/8\pi}}\right) \cdot \sqrt{\frac{(\rho_{rel})_0 a_0^4}{a^4(t^*)} + \frac{(\rho_m)_0 a_0^3}{a^3(t^*)}} \sim 10^{-45} \cdot 10^1 \cdot \sqrt{10^{80}} \approx 10^{-4} << 1 \tag{9}$$

If we examine what happens with $|\Lambda_{5-\text{dim}}| \sim c_2 T^{-\beta}$

TABLE 1

Cosmological Λ in 5 and 4 dimensions [5]

Time $0 \le t << t_P$	**Time** $0 \le t < t_P$	**Time** $t \ge t_P$	**Time** $t > t_P \rightarrow$ **today**
$\lvert\Lambda_5\rvert$ undefined, $T \approx \varepsilon^+ \rightarrow T \approx 10^{32} K$ $\Lambda_{4-\text{dim}} \approx$ almost ∞	$\lvert\Lambda_5\rvert \approx \varepsilon^+$, $\Lambda_{4-\text{dim}} \approx$ extremely large $10^{32} K > T > 10^{12} K$	$\lvert\Lambda_5\rvert \approx \Lambda_{4-\text{dim}}$, T much smaller than $T \approx 10^{12} K$	$\lvert\Lambda_5\rvert \approx$ huge, $\Lambda_{4-\text{dim}} \approx$ constant, $T \approx 3.2 K$

For times $t > t_P \rightarrow$ today, a stable instanton is assumed, along the lines brought up by t'Hooft [8], due to the stable $\Lambda_{4-\dim} \approx$ constant ~ very small value, roughly at the value given today. This assumes a radical drop-off of the cosmological constant for, say right after the electroweak transition. This would be in line with Kolb's assertion of the net degrees of freedom in space-time drop from about 1000 to less than two, especially if $t > t_P \rightarrow$ today in terms of the value of time after the big bang. The supposition we are making here is that the value of N so obtained is actually proportional to a numerical graviton density we will refer to as <n>., provided that there is a bias toward HFGW, which would mandate a very small value for $V \approx R_H^3 \approx \lambda^3$.Furthermore, structure formation arguments, as given by Perkins [9] give ample evidence that if we use an energy scale, m, over a Planck mass value M_{Planck}, as well as contributions from field amplitude ϕ, and using the contribution of scale factor behavior $\frac{\dot{a}}{a} \equiv H \approx -m \cdot \frac{\phi}{3 \cdot \dot{\phi}}$, where we assume $\ddot{\phi} \cong 0$ due to inflation

$$\frac{\Delta\rho}{\rho} \sim H\Delta t \sim \frac{H^2}{\dot{\phi}} \sim \left(\frac{m}{M_{Planck}}\right) \times \left(\frac{\phi}{M_{Planck}}\right) \sim 10^{-5} \tag{10}$$

At the very onset of inflation, $\phi << M_{Planck}$, and if m (assuming $\hbar = c = 1$) is due to inputs from a prior universe, we have a wide range of parameter space as to ascertain where $\Delta S \approx \Delta N_{gravitons} \neq 10^{88}$ [9]comes from and plays a role as to the development of entropy in cosmological evolution In the next Chapter , we will discuss if or not it is feasible / reasonable to have data compression of prior universe 'information'. It suffices to say that if $S_{initial} \sim 10^5$ is transferred from a prior universe to our own universe at the onset of inflation,, at times less than Planck time $t_P \sim 10^{-44}$ seconds, that enough information **MAY** exit for the preservation of the prior universe's cosmological constants, i.e. $\hbar, G, \alpha$ (fine structure constant) and the like. We do not have a reference for this and this supposition is being presented for the first time. Times after after t= 10^-44 are not less important. But that the 'constant's memory' is already imprinted in the universe, so to speak. I.e. a memory transfer is implied as far as being transferred from the beginning. Confirmation of this hypothesis depends upon models of how much 'information' $\hbar, G, \alpha$ actually require to be set in place, at the onset of our universe's inflation, a topic which we currently have no experimental way of testing at this current time. Issues raised in [10], [11], [12], [13] are important as to the research protocols

Consider now what could happen with a phenomenological model bases upon the following inflection point i.e. split regime of different potential behavior

$$V(\phi) = g \cdot \phi^{\alpha} \tag{13}$$

De facto, what we come up with pre, and post Planckian space time regimes, when looking at consistency of the emergent structure is the following. Namely by addusting what is done by Weinberg[14] we have [15],

$$V(\varphi) \propto \phi^{|\alpha|} \qquad \text{for } t < t_{PLanck} \tag{14}$$

Also, we would have

$$V(\varphi) \propto 1/\phi^{|\alpha|} \qquad \text{for } t >> t_{PLanck} \tag{15}$$

The switch between Eq. (14) and Eq. (15) is not justified analytically. I.e. it breaks down. Beckwith et al (2011) designated this as the boundary of a causal discontinuity. Now according to Weinberg [14] , if

$\in = \frac{\lambda^2}{16\pi G}, H = 1/\in t$ so that one has a scale factor behaving as [15]

$$a(t) \propto t^{1/\in} \tag{16}$$

Then, if [14]

$$|V(\phi)| << (4\pi G)^{-2} \tag{17}$$

there are no quantum gravity effects worth speaking of. I.e., if one uses an exponential potential a scalar field could take the value of , when there is a drop in a field from ϕ_1 to ϕ_2 for flat space geometry and times t_1 to t_2 [15]

$$\phi(t) = \frac{1}{\lambda}\ln\left[\frac{8\pi Gg \in^2 t^2}{3}\right] \tag{18}$$

Then the scale factors, from Planckian time scale as [15]

$$\frac{a(t_2)}{a(t_1)} = \left(\frac{t_2}{t_1}\right)^{1/\in} = \exp\left[\frac{(\phi_2 - \phi_1)\lambda}{2\in}\right] \tag{19}$$

The more $\frac{a(t_2)}{a(t_1)} >> 1$, then the less likely there is a tie in with quantum gravity. Note those that the way this potential is defined is for a flat , Roberson-Walker geometry, and that if and when $t_1 < t_{Planck}$ then what is done in Eq. (11) no longer applies, and that one is no longer having any connection with even an octonionic Gravity regime.

Increase in degrees of freedom in the sub Planckian regime.

Starting with [16] , [17]

$$E_{thermal} \approx \frac{1}{2} k_B T_{temperature} \propto [\Omega_0 \breve{T}] \sim \tilde{\beta} \tag{21}$$

The assumption is that there would be an initial fixed entropy arising, with $\overline{N}$ as a nucleated structure arising in a short time interval as a temperature $T_{temperature}\varepsilon(0^+, 10^{19}GeV)$ arrives. One then obtains, dimensionally speaking [16], [17]

$$\frac{\Delta\tilde{\beta}}{dist} \cong (5k_B \Delta T_{temp}/2) \cdot \frac{\overline{N}}{dist} \sim qE_{net-electric-field} \sim [T\Delta S/dist] \tag{22}$$

The parameter, as given by $\Delta\tilde{\beta}$ will be one of the parameters used to define chaotic Gaussian mappings. Candidates as to the inflation potential would be in powers of the inflation, i.e. in terms of ϕ^N , with N=4 effectively ruled out, and perhaps N=2 an admissible candidate (chaotic inflation). For N = 2, one gets [16], [16]

$$[\Delta S] = [\hbar/T] \cdot \left[2k^2 - \frac{1}{\eta^2}\left[M^2_{Planck} \cdot \left[\left[\frac{6}{4\pi} - \frac{12}{4\pi}\right] \cdot \left[\frac{1}{\phi}\right]^2 - \frac{6}{4\pi} \cdot \left[\frac{1}{\phi^2}\right]\right]\right]\right]^{1/2} \sim n_{Particle-Count} \tag{23}$$

If the inputs into the inflation, as given by ϕ^2 becomes from Eq. (6) a random influx of thermal energy from temperature, we will see the particle count on the right hand side of Eq. (23) above a partly random creation of $n_{Particle-Count}$ which we claim has its counterpart in the following treatment of an increase in degrees of freedom. The way to introduce the expansion of the degrees of freedom from nearly zero, at the maximum point of contraction to having **N(T)~ 10³** is to first define the classical and quantum regimes of gravity in such a way as to minimize the point of the bifurcation diagram affected by quantum processes.[16] If we suppose smoothness of space time structure down to a grid size of $l_{Planck} \sim 10^{-33}$ centimeters at the start of inflationary expansion we have when doing this construction what would be needed to look at the maximum point of contraction, setting at $l_{Planck} \sim 10^{-33}$ centimeters the quantum 'dot' or infometron, as a de facto measure zero set, as the bounce point, with classical physics behavior before and after the bounce 'through' the quantum dot. Dynamical systems modeling could be directly employed right 'after' evolution through the 'quantum dot' regime, with a transfer of crunched in energy to Helmholtz free energy, as the driver 'force' for a Gauss map type chaotic diagram right after the transition to the quantum 'dot' point of maximum contraction. The diagram, in a bifurcation sense would look like an application of the Gauss mapping of [16].[17]

$$x_{i+1} = \exp\left[-\tilde{\alpha} \cdot x_i^2\right] + \tilde{\beta} \tag{24}$$

In dynamical systems type parlance, one would achieve a diagram, with tree structure looking like what was given by Binous [18], using material written up by Lynch [19] .Now that we have a model as to what could be a change in space time geometry, let us consider what may happen during the Higgs mechanism and why it may not apply as expected in very early universe geometry

Higgs Mechanism, and its consequence in the onset of inflation. I.e. why it could break down

Let us begin first with a U(1) gauge theory, the Fermion ψ would transform locally as given by [20]

$$\psi \to \psi' = \left(\exp\left[-ig\vartheta \cdot q(x)\right]\right) \cdot \psi \tag{25}$$

This has a Lagrangian given by, an expression for covariant derivative $D_\mu = \partial_\mu + ig\vartheta A_\mu(x)$, and also $F_{\mu\nu} = \partial_\mu A_\nu - \partial_\nu A_\mu$, so that $\partial_\mu F_{\mu\nu} = j_\nu$ for current. With the mass term for the gauge boson A_μ not allowed by gauge symmetry via the Lagrangian $\varsigma = i\overline{\psi}\gamma^\mu D_\mu \psi + m\overline{\psi}\psi - \frac{1}{4}F^{\mu\nu}F_{\mu\nu}$

A way to allow for the mass to be factored in, i.e. look at $\phi \to \phi' = \left(\exp\left[-ig\vartheta \cdot q(x)\right]\right) \cdot \phi$, and then

$$\varsigma(\phi) = iD^{\varpi}\phi^+ D_\mu \phi - \frac{1}{2}\mu^2 \phi^+ \phi - \frac{1}{4}\lambda\left(\phi^+\phi\right)^2 \tag{26}$$

If $\mu^2 < 0$, the potential has a minimum, with $\langle \phi^+ \phi \rangle = v^2 = -\mu^2 / \lambda > 0$, with a VeV $\langle \phi \rangle = v$. Then

$$\phi = (\eta + v)\exp\left[i\sigma / v\right] \tag{27}$$

As stated by U. Sarkar [20], a kinetic energy term for the scalar field, namely $g^2 v^2 A^\mu A_\mu \subset D^\mu \phi^+ D_\mu \phi$ is such that a mass term may exist. Now as to why it is stated that this procedure may break down. A scalar field will no longer be massless if the following step is taken, namely an explicit symmetry breaking term $m^2\left(\phi\phi + \phi^*\phi^*\right)$ will allow a scalar field ϕ to be expanded about a VeV $\langle \phi \rangle = v$ with

$$\phi = (\eta + v)\exp\left[i\sigma / v\right] \sim \eta + v + i\sigma - \sigma^2 / 2v \tag{28}$$

so that the mass of σ is m^2, so σ is a pseudo nambu goldstone boson. If one wishes to have explicit examples of the VeVs, then consider [20]

$$SU(5) \to SU(4) \times U(1) \Rightarrow \langle \phi \rangle = \begin{pmatrix} 1 & 0 & 0 & 0 & 0 \\ 0 & 1 & 0 & 0 & 0 \\ 0 & 0 & 1 & 0 & 0 \\ 0 & 0 & 0 & 1 & 0 \\ 0 & 0 & 0 & 0 & -4 \end{pmatrix} \tag{29}$$

$$SU(5) \rightarrow SU(3)\times SU(2)\times U(1) \Rightarrow \langle \phi \rangle = \begin{pmatrix} 1 & 0 & 0 & 0 & 0 \\ 0 & 1 & 0 & 0 & 0 \\ 0 & 0 & 1 & 0 & 0 \\ 0 & 0 & 0 & -3/2 & 0 \\ 0 & 0 & 0 & 0 & -3/2 \end{pmatrix} \tag{30}$$

In the case of when one is looking at when the VeV is congruent with a broken symmetry potential, as of the form $m^2(\phi\phi + \phi^*\phi^*)$, which no longer exists in the situation where one is looking at k essence inflation, we then will be having to consider the situation is given by: The main point as to why the Higgs paradigm may break down lies in the fact that emergent structure can be formulated without use of a broken symmetry potential as given by $m^2(\phi\phi + \phi^*\phi^*)$.

How to have particle formation without a broken symmetry potential. Use of Sherrer k Esesence

. In particular, the situation to watch can be diagrammed out [21]by appendix entry where we are looking at the k essence scenario . This means we have a small value for the 'growth of density perturbations' [20], [22]

$$C_S^2 \cong \frac{1}{1+2\cdot(X_0+\tilde{\varepsilon}_0)\cdot(1/\tilde{\varepsilon}_0)} \equiv \frac{1}{1+2\cdot\left(1+\frac{X_0}{\cdot\tilde{\varepsilon}_0}\right)} \tag{31}$$

if we can approximate

$$(\partial_\mu \phi)\cdot(\partial^\mu \phi) \equiv \left(\frac{1}{c}\cdot\frac{\partial\phi}{\partial\cdot t}\right)^2 - (\nabla\phi)^2 \cong -(\nabla\phi)^2 \rightarrow -\left(\frac{d}{dx}\phi\right)^2 \tag{31a}$$

a comparatively small contribution w.r.t. time variation, but a very large in many cases contribution w.r.t.

spatial variation of phase

$$|X_0| \approx \frac{1}{2}\cdot\left(\frac{\partial\phi}{\partial x}\right)^2 >> \tilde{\varepsilon}_0 \tag{31b}$$

$$0 \le C_S^2 \approx \varepsilon^+ << 1 \tag{32}$$

and

$$w \equiv \frac{p}{\rho} \cong \frac{-1}{1-4\cdot(X_0+\tilde{\varepsilon}_0)\cdot\left(\frac{F_2}{F_0+F_2\cdot(\tilde{\varepsilon}_0)^2}\cdot\tilde{\varepsilon}_0\right)} \approx 0 \tag{33}$$

We get these values for the phase being nearly a 'box' of height approximately scaled to be about $2\cdot\pi$ and of width ***L*** . Which we obtained by setting [23]

$$\phi \approx \pi \cdot \left[\tanh b\cdot(x+L/2) - \tanh b\cdot(x-L/2)\right] \tag{34}$$

This means that the initial conditions we are hypothesizing are in line with the equation of state conditions appropriate for a cosmological constant but near zero effective sound speed . As it is, we approximate

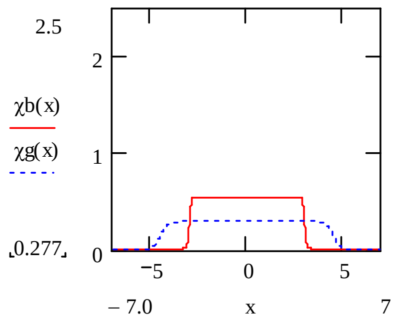

Fig 1

Evolution of the phase from a thin wall approximation to a more nuanced thicker wall approximation with increasing L between S-S' instanton componets. The 'height' drops and the 'width' L increases corresponds to a de evolution of the thin wall approximation. This is in tandem with a collapse of an initial nucleating 'potential' system to the standard chaotic scalar ϕ^2 *potential system of Guth***[24]..** *As the 'hill' flattens, and the thin wall approximation dissipates, the physical system approaches standard cosmological constant behavior.*

This is occurring in the regime in which Octonian gravity initially does not apply and which eventually it does apply. So, let us look at the following

Relevance to Octonian Quantum gravity constructions? Where does non commutative geometry come into play?

Crowell [24] wrote on page 309 that in his Eq. (8.141), namely

$$\left[x_j, p_i\right] \cong -\beta\cdot\left(l_{Planck}/l\right)\cdot\hbar T_{ijk}x_k \to i\hbar\delta_{i,j} \tag{36}$$

Here, β is a scaling factor, while we have, above, after a certain spatial distance, a Kroniker function so that at a small distance from the confines of Planck time, we recover our quantum mechanical behavior. Our contention is, that since Eq. (26) depends upon Energy- momentum being conserved as an average about quantum fluctuations, that if energy-momentum is violated, in part, that Eq. (36) falls apart. How Crowell forms Eq. (36) at the Planck scale depends heavily upon Energy- Momentum being conserved.[25] Our construction VIOLATES energy – momentum conservation. N. Poplawski[26], [27] also has a very revealing construction for the vacuum energy, and cosmological constant which we reproduce, here

$$\Lambda = \left[\frac{3\kappa^2}{16}\right]\cdot\left(\bar{\psi}\gamma_j\gamma^5\psi\right)\cdot\left(\bar{\psi}\gamma^j\gamma^5\psi\right) \text{ And } \rho_\Lambda = \left[\frac{3\kappa}{16}\right]\cdot\left(\bar{\psi}\gamma_j\gamma^5\psi\right)\cdot\left(\bar{\psi}\gamma^j\gamma^5\psi\right) \tag{37}$$

Poplawski writes that formation of the above, is:

"Such a torsion-induced cosmological constant depends on spinor fields, so it is not constant in time (it is constant in space at cosmological scales in a homogeneous and isotropic universe). However, if

these fields can form a condensate then the vacuum expectation value of (**Eq. 37**) *will behave like a real cosmological constant"*

Poplawski [25],[26]write his formulation in terms of a quark- gluon QCD based condensate. Our contention is that once a QCD style condensate breaks up there will afterwards be NO equivalent structure to Eq. (37) and Eq. (38) even at the beginning of inflation right after the break down of space time particle transfer .Once that condensate structure is not possible then as quantified by Eq. (8.140) of Crowell [25], the following will not hold:

$$\oint p_i \, dx_k = \hbar \delta_{i,k} \tag{39}$$

Eq. (8.40) of the Crowell [25]manuscript also makes the additional assumption, that non flat space has a geometric non-commutativity protocol which is delineated by the following spatial relationship. When Eq. (40) goes to zero, we recover the regime in which quantum mechanics holds.

$$\left[x_j, x_k\right] = \beta \cdot l_P \cdot T_{j,k,l} \cdot x_l \tag{40}$$

Does the (QCD) condensate occur post plankian, and not work for pre plankian regime ? Yes. The problem lies with Eq. (8.140) of Crowell [25] with the final equality not holding. If one were integrating across a causal barrier,

$$\oint [x_j, p_i] dx_k \approx -\oint p_i [x_j, dx_k] = -\beta \cdot l_P \cdot T_{j,k,l} \oint p_i dx_l \neq -\hbar\beta \cdot l_P \cdot T_{i,j,k} \tag{41}$$

Very likely, across a causal boundary, between $\pm l_P$ across the boundary due to the causal barrier, one would have

$$\oint p_i \, dx_k \neq \hbar \delta_{i,k}, \oint p_i \, dx_k \equiv 0 \tag{42}$$

I.e.

$$\left. \oint_{\pm l_P} p_i dx_k \right|_{i=k} \to 0 \tag{43}$$

If so, then [25]

$$\left[x_j, p_i\right] \neq -\beta \cdot \left(l_{Planck} / l\right) \cdot \hbar T_{ijk} x_k \quad and \quad does \quad not \to i\hbar \delta_{i,j} \tag{44}$$

Eq. (44) in itself would mean that in the pre Planckian physics regime, and in between $\pm l_P$, QM no longer applies. What we will do next is to begin the process of determining a regime in which Eq. (34) may no longer hold via experimental data sets. As an example of present confusion, please consider the following discussion where leading cosmologists, i.e. Sean Carroll [28](2005) asserted that there is a distinct possibility that mega black holes in the center of spiral galaxies have more entropy, in a calculated sense, i.e. up to 10^{90} in non dimensional units. This has to be compared to Carroll's (2005)[28] stated value of up to 10^{88} in non dimensional units for observable non dimensional entropy units for the observable universe. Assume that there are over one billion spiral galaxies, with massive black holes in their center, each with entropy 10^{90}, and then there is due to spiral galaxy entropy contributions $10^6 \times 10^{90} = 10^{96}$ entropy units to contend with, vs. 10^{88} entropy units to contend with for the observed universe. I.e. at least a ten to the eight order difference in entropy magnitude to contend with. A further datum to consider is that Eq. (44) with its variance of density fluctuations may eventually be linkable to Kolmogrov theory as far as structure formation . If we look at R. M. S. Rosa [29] (2006) , and energy cascades of the form of the 'energy dissipation law' , assuming u_0, l_0 are minimum velocity and length, with velocity less than the speed of light, and the length at least as large, up to 10^6 time larger than Planck length l_{Planck}

$$\varepsilon \approx \frac{u_0^3}{l_0} \tag{45}$$

Eq. (45) above can be linked to an eddy break down process, which leads to energy dissipated by viscosity. If applied appropriately to structures transmitted through a 'worm hole' from a prior to a present universe, it can explain

1) How there could be a break up of ‘encapsulating’ structure which may initially suppress additional entropy beyond $S_{initial} \sim 10^{5}$, in the onset of inflation
2) Provide a ‘release’ mechanism $\Delta S \approx \Delta N_{gravitons} < 10^{54} << 10^{88}$, with $\Delta S \approx \Delta N_{gravitons} \sim 10^{21}$ perhaps a starting point for increase in entropy in $\Delta t \approx t_{Planck} \sim 5 \times 10^{-44}$ sec, rising to $\Delta S \approx \Delta N_{gravitons} \leq 10^{54} << 10^{88}$ for times up to 1000 seconds after the big bang.

Let us now consider the impact of the octonian gravity paradigm and where it may break down. And why.

Finally, Relic graviton produced entropy at the onset of the big bang . Why starting entropy would be so small while CMBR entropy would be so large

As a closing remark, Beckwith wishes to suggest a solution to Penrose’s implied question about entropy as raised in Edingborough , Scotland [31] conference proceedings. Penrose talks about the 2nd law, and its implied requirements as to the small initial value of early universe entropy, and then states that gravitational entropy would not be so major, whereas CMBR matter contributed entropy would be much larger. Beckwith is convinced that relic graviton production at the onset of the big bang, i.e. before the contribution of entropy from matter itself would be necessary to boost entropy from its small 10^{5} value at the onset of the big bang, to a much higher level, and that entropy would be initially dramatically boosted by that process. I.e. the uniformity requirement Penrose talks about in structure would be actually as of up to the Electro weak transition, and far after the initial onset of inflation itself.

A new idea extending Penrose’s suggestion of cyclic universes, black hole evaporation, and the embedding structure our universe is contained within

Beckwith strongly suspects that there are no fewer than N (a large number) of universes under going Penrose ‘infinite expansion’ and all these are contained within a mega universe structure. Furthermore, that each of the N universes has black hole evaporation commencing, with the Hawking radiation from decaying black holes. If each of the N universes is definable by a partition function, we can call $\left\{\Xi_i\right\}_{i \equiv N}^{i \equiv 1}$, then there exist an information minimum ensemble of mixed minimum information roughly correlated as about $10^{7} - 10^{8}$ bits of information per each partition function in the set $\left.\left\{\Xi_i\right\}_{i \equiv N}^{i \equiv 1}\right|_{before}$, so minimum information is conserved between a set of partition functions per each universe

$$\left.\left\{\Xi_i\right\}_{i \equiv N}^{i \equiv 1}\right|_{before} \equiv \left.\left\{\Xi_i\right\}_{i \equiv N}^{i \equiv 1}\right|_{after} \tag{46}$$

However, that there is non uniqueness of information put into each partition function $\left\{\Xi_i\right\}_{i \equiv N}^{i \equiv 1}$. Furthermore that within the mega structure, that Hawking radiation from the black holes is collated via a strange attractor collection in the mega universe structure to form a new big bang for each of the N universes as represented by $\left\{\Xi_i\right\}_{i \equiv N}^{i \equiv 1}$. Verification of this mega structure compression and expansion of information with a non unique venue of information placed in each of the N universes would strongly favor Ergodic mixing treatments of initial values for each of the N universes expanding from a quasi singularity beginning. If this idea is in any way confirmable, it would lend credence as to the formation of the dark flow hypothesis, and of how anharmonic perturbative contributions to initial inflationary expansion may occur , within a partially random ergotic background. Beckwith claims that such a process would inherently favor the small 10^{7} bits of information per each partition function representing the ‘start’ of expansion of a new universe. Hopefully, in doing so, one can explain, energy flux being re formulated for each universe.

I.e. start with the Alcubierre's formalism about energy flux, assuming that there is a solid angle for energy distribution Ω for the energy flux to travel through. [32]

$$\frac{dE}{dt} = [\lim r \to \infty]\left[\frac{r^2}{16\pi}\right]\oint\left|\int_{-\infty}^{t}\Psi_4 dt'\right|^2 \cdot d\Omega \tag{47}$$

The expression Ψ_4 is a Weyl scalar which we will, before the electro weak phase transition, assume that time dependence of both h^+ and h^x is miniscule and that initially $h^+ \approx h^x$, so as to initiate Ψ_4 as

$$\Psi_4 \cong -\frac{1}{4}\cdot\left[+\partial_r^2 h^+\right]\cdot(-1+i) \tag{48}$$

The upshot, is that the initial energy flux about the inflationary regime would lead to looking at[31],[32]

$$\left|\int_{-\infty}^{t}\Psi_4 dt'\right| \approx \left|\frac{1}{2}\cdot\left[+\partial_r^2 h^+\right]\right|\cdot\left(\tilde{n}\cdot t_{Planck}\right) \tag{49}$$

This will lead to an initial changing energy flux at the onset of inflation which will be presented as

$$\frac{dE}{dt} = \left[\frac{r^2}{64\pi}\right]\cdot\left|+\partial_r^2 h^+\right|^2\cdot\left[\tilde{n}\cdot t_{Planck}\right]^2\cdot\Omega \tag{50}$$

If we are talking about an initial energy flux, we then can approximate the above as[31],[32]

$$E_{initial-flux} \cong \left[\frac{r^2}{64\pi}\right]\cdot\left|+\partial_r^2 h^+\right|^2\cdot\left[\tilde{n}\cdot t_{Planck}\right]^3\cdot\Omega_{effective} \tag{51}$$

Inputs into both the expression $\left|\partial_r^2 h^+\right|$, as well as $\Omega_{effective}$ will comprise the rest of this document, plus our conclusions. The derived value of $\Omega_{effective}$ as well as $E_{initial-flux}$ will be tied into a way to present energy per graviton, as a way of obtaining n_f. The n_f value so obtained, will be used to make a relationship , using Y. J. Ng's entropy [9] counting algorithm of roughly [10]. $S_{entropy} \sim n_f$. We assert that in order to obtain $S_{entropy} \sim n_f$ from initial graviton production, as a way to quantify n_f, that a small mass of the graviton can be assumed. How to tie in this energy expression , as given in Eq. (51) will be to look at the formation of a non trivial gravitational measure which we can state as a new big bang for each of the N universes as represented by [31],**[32]** and $n(E_i)\cdot$ the density of states at a given energy E_i for a partition function defined by **[11] , [31],[32]**

$$\{\Xi_i\}_{i\equiv 1}^{i\equiv N} \propto \left\{\int_0^\infty dE_i\cdot n(E_i)\cdot e^{-E_i}\right\}_{i\equiv 1}^{i\equiv N}. \tag{52}$$

Each of the terms E_i would be identified with Eq.(52) above, with the following iteration given, namely for N universes

$$\frac{1}{N}\cdot\sum_{j=1}^{N}\Xi_j\Big|_{j-before-nucleation-regime}\xrightarrow[vacuum-nucleation-tranfer]{}\Xi_i\Big|_{i-fixed-after-nucleation-regime} \tag{53}$$

For N number of universes, with each $\Xi_j\Big|_{j-before-nucleation-regime}$ for j = 1 to N being the partition function of each universe just before the blend into the RHS of Eq. (54) above for our present universe. Also, each of the independent universes given by $\Xi_j\big|_{j-before-nucleation-regime}$ would be constructed by the absorption of one million black holes sucking in energy**. I.e. in the end**

$$\Xi_j\big|_{j-before-nucleation-regime} \approx \sum_{k=1}^{Max} \tilde{\tilde{\Xi}}_k\big|_{black-holes-jth-universe} \tag{54}$$

One can treat Eq. (54) as a de facto Ergodic mixing of prior universes to a present universe, with the partition function of each of the universes defined by Eq. (53) above. Filling in the inputs into Eq. (52) to Eq. (54) is what will be done in the months ahead. $\left|\partial_r^2 h^+\right|$ will be the one to fill in, via considering [31] plus other models. Doing so will begin to allow us to form more precise evaluations of Eq. (52) to Eq. (54) . Making sense of $\left|\partial_r^2 h^+\right| \sim k^2 h^+$ requires that we understand the evolution of gravity waves and gravitons as a k essence phenomenon. This is part of our future works

Summary as to what is known, and not known about the Null Energy Condition in Cosmology . And information exchange between Prior to Present Universes.

As stated in [1], the NEC is linked to the following, ie. Look at the general null energy condition first

The null energy condition stipulates that for every future-pointing null vector field (for all of the GR) $\vec{k}$,

- $$\rho = T_{ab}\, k^a\, k^b \geq 0. \tag{55}$$

With respect to a frame aligned with the motion of the matter particles, the components of the matter tensor take the diagonal form, in Euclidian space that

$$T^{\hat{a}\hat{b}} = \begin{bmatrix} \rho & 0 & 0 & 0 \\ 0 & p & 0 & 0 \\ 0 & 0 & p & 0 \\ 0 & 0 & 0 & p \end{bmatrix}. \tag{56}$$

The simplest statement of the Null energy condition is that he null energy condition stipulates that

$$\rho + p \geq 0. \tag{57}$$

I.e. the equation of state to consider is, if $w \leq -1$, then if what [1] suggests is true, then there will be a reason to consider the relative import of Eq. (1), Eq. (2), and Eq. (4) in terms of contributions. I.e. we do have problems with the idea of variance of the cosmological constant, G, and will reference it. We also will build upon the consequences of $w \leq -1$, and Eq. (1), Eq. (2), and Eq. (4) as far as first principle treatments of if information exchange between a prior to present universe is feasible and experimentally testable.

One of the simplest examples of a break down of the NEC, is the Casmir effect. I.e. it is topological , in that the sign of the vacuum energy depends on both the geometry and topology of the configuration For what it is worth, the Casmir effect is between two flat plates. We can generalize this idea to initial domain wall physics . Spherical geometry as we know it does not violate the NEC. Further domain wall physics may lead to a break down of the NEC.

We also refer to a treatment of the NEC ,and its possible consequences if we look at an effective Friedman equation as given by [15], as seen by

$$H_a^2 = \left[\frac{\dot{a}}{a}\right]^2 = \frac{8\pi G}{3} G \cdot \rho + \left(\frac{\kappa^2}{24 \cdot (3+n)}\right) \cdot \left[2 + n \cdot \left[1 - 3 \cdot (\nu - w)\right]\right] \cdot \rho^2 - e \qquad (58)$$

The scaling done in this situation has [15], especially if e is a constant in Eq. (58)

$$\rho = a^{-3 \cdot [1+w]} \qquad (59)$$

As stated in [16]. We expect that there will be flat space geometry almost in the beginning of the early big bang. I.e. this will lead to Eq(59) , if $w < -1$ implying that $\rho = a^{-3 \cdot [1+w]} \sim a^{+\varepsilon} \to 0^{+}$ if there is a violation of the NEC.

Now let us review, briefly, a few items as far as information exchange , and what to consider. To do this, look at the following. As quoted from [35]. We consider t he inter play of bits to information. I.e. as seen in a colloquium presentation done by Dr. Smoot in Paris [35] (2007); he alluded to the following information theory constructions which bear consideration as to how much is transferred between a prior to the present universe in terms of information 'bits'.

0) Physically observable bits of information possibly in present Universe - 10^{180}
1) Holographic principle allowed states in the evolution / development of the Universe - 10^{120}
2) Initially available states given to us to work with at the onset of the inflationary era- 10^{10}
3) Observable bits of information present due to quantum / statistical fluctuations -10^{8}

Our guess is as follows. That the thermal flux accounts for perhaps 10^{10} bits of information. These could be transferred from a prior universe to our present , and that there could be , perhaps 10^{120} minus 10^{10} bytes of information temporarily suppressed during the initial bozonification phase of matter right at the onset of the big bang itself .

Beckwith [36] stated a criteria as far as graviton production, and a toy model of the universe. If one has Eq. [2] shut off due to $w < -1$, so then that $\rho = a^{-3\cdot[1+w]} \sim a^{+\varepsilon} \to 0^{+}$ occurs, then the causal discontinuity so references in [35], [37] by Beckwith et al, will have major consequences as far as a away to determine if gravitons have a small mass, and if there is a way to determine if a prior universe has contribution as to the information transferred as to the present universe. We will now assume, that the catastrophe given as stated by $\rho = a^{-3\cdot[1+w]} \sim a^{+\varepsilon} \to 0^{+}$ does not occur. We will then refer to how one could have

First principles argument as to large scale values of the absolute magnitude of the cosmological vacuum energy

Look at an argument provided by Thanu Padmanabhan [38], leading to the observed cosmological constant value suggested by Park. Assume that $l_P \sim 10^{-33} cm \xrightarrow[Quantum-Gravity-threshold]{} \tilde{N}^{\alpha} \cdot l_P$, but that when we make this substitution that $1 \le \tilde{N}^{\alpha} \le 10^{2}$ [39]

$$\begin{aligned} \rho_{VAC} &\sim \frac{\Lambda_{observed}}{8\pi G} \sim \sqrt{\rho_{UV} \cdot \rho_{IR}} \\ &\sim \sqrt{l_{Planck}^{-4} \cdot l_{H}^{-4}} \sim l_{Planck}^{-2} \cdot H_{observed}^{2} \end{aligned} \tag{60}$$

$$\Delta\rho \approx \text{a dark energy density} \sim H_{observed}^{2} / G \tag{61}$$

We can replace $\Lambda_{observed}, H_{observed}^{2}$ by $\Lambda_{initial}, H_{initial}^{2}$. In addition we may look at inputs from the initial value of the Hubble parameter to get the necessary e folding needed for inflation, according to

$$\begin{aligned} &E - foldings = H_{initial} \cdot \left(t_{End\ of\ \inf} - t_{beginning\ \ of\ \inf} \right) \equiv N \ge 100 \\ &\Rightarrow H_{initial} \ge 10^{39} - 10^{43} \end{aligned} \tag{62}$$

Leading to

$$a(End - of - \inf) / a(Beginning - of - \inf) \equiv \exp(N) \tag{63}$$

If we set $\Lambda_{initial} \sim c_1 \cdot \left[T \sim 10^{32}\, Kelvin \right]$ implying a very large initial cosmological constant value, we get in line with what Park suggested for times much less than the Planck interval of time at the instant of nucleation of a vacuum state

$$\Lambda_{initial} \sim \left[10^{156} \right] \cdot 8\pi G \approx huge \quad number \tag{64}$$

It is easy to infer, with minimum effort that Eq. (2) and Eq. (64) give much the same information. Provided that $w < -1$, our argument that inflation needs Eq. (2) is confirmation as to what was said in [1]. If we avoid then, having $w < -1$, then the following may hold and needs experimental verification.

Minimum amount of information needed to initiate placing values of fundamental cosmological parameters

A.K. Avessian's [40] article (2009) about alleged time variation of Planck's constant from the early universe depends heavily upon initial starting points for $\hbar(t)$, as given below, where we pick our own values for the time parameters, for reasons we will justify in this manuscript:

$$\hbar(t) \equiv \hbar_{initial}\left[t_{initial} \le t_{Planck}\right] \cdot \exp\left[-H_{macro} \cdot \left(\Delta t \sim t_{Planck}\right)\right] \tag{65}$$

The idea is that we are assuming a granular , discrete nature of space time. Futhermore, after a time we will state as $t \sim t_{Planck}$ there is a transition to a present value of space time, which is then probably going to be held constant.

It is easy to, in this situation, to get an inter relationship of what $\hbar(t)$ is with respect to the other physical parameters , i.e. having the values of α written as $\alpha(t) = e^2 / \hbar(t) \cdot c$, as well as note how little the fine structure constant actually varies . Note that if we assume an unchanging Planck's mass $m_{Planck} = \sqrt{\hbar(t) c / G(t)} \sim 1.2 \times 10^{19}\, GeV$, this means that G has a time variance, too.

This leads to us asking what can be done to get a starting value of $\hbar_{initial}\left[t_{initial} \le t_{Planck}\right]$ recycled from a prior universe, to our present universe value. What is the initial value, and how does one insure its existence?

We obtain a minimum value as far as 'information' via appealing to Hogans [41] (2002) argument where we have a maximum entropy as

$$S_{\max} = \pi / H^2 \tag{66}$$

, and this can be compared with A.K. Avessian's article [40] (2009) value of, where we pick $\Lambda \sim 1$

$$H_{macro} \equiv \Lambda \cdot \left[H_{Hubble} = H\right] \tag{67}$$

I.e. a choice as to how $\hbar(t)$ has an initial value, and entropy as scale valued by $S_{\max} = \pi / H^2$ gives us a ball park estimate as to compressed values of $\hbar_{initial}\left[t_{initial} \le t_{Planck}\right]$ which would be transferred from a prior universe, to todays universe. If $S_{\max} = \pi / H^2 \sim 10^5$, this would mean an incredibly small value for the INITIAL H parameter, i.e. in pre inflation, we would have practically NO increase in expansion, just before the introduction vacuum energy, or emergent field energy from a prior universe, to our present universe.

Typically though, the value of the Hubble parameter, during inflation itself is HUGE, i.e. H is many times larger than 1, leading to initially very small entropy values. This means that we have to assume, initially, for a minimum transfer of entropy/ information from a prior universe, that H is neligible. If we look at Hogan's holographic model, this is consistent with a non finite event horizon [41]

$$r_0 = H^{-1} \tag{68}$$

This is tied in with a temperature as given by

$$T_{black-hole} = (2\pi \cdot r_0)^{-1} \tag{69}$$

Nearly infinite temperatures are associated with tiny event horizon values, which in turn are linked to huge Hubble parameters of expansion. Whereas initially nearly zero values of temperature can be arguably linked to nearly non existent H values, which in term would be consistent with $S_{\max} = \pi / H^2 \sim 10^5$ as a starting point to entropy. We next then must consider how the values of initial entropy are linkable to other physical models. I.e. can there be a transfer of entropy/ information from a pre inflation state to the present universe. Doing this will require that we keep in mind, as Hogan writes, that the number of distinguishable states is writable as [41]

$$N = \exp(\pi H^{-2}) \tag{70}$$

If , in this situation, that N is proportional to entropy, i.e. N as ~ number of entropy states to consider, , then as H drops in size, as would happen in pre inflation conditions, we will have opportunities for N ~ 10^5

1st level of Conclusion: Several reasons for the Analog nature of reality with digital a sub set of a larger Analog basis

We wish to summarize what we have presented in an orderly fashion. Doing so is a way of stating that Analog, reality is the driving force behind the evolution of inflationary physics

a) Pre Octonian gravity physics (analog regime of reality) featurea a break down of the Octonian gravity commutation relationships when one has curved space time. **This corresponds, as brought up in the Jacobi iterated mapping for the evolution of degrees of freedom to a build up of temperature as themal heat influx for an increase in degrees of freedom from 2 to over 1000. Per unit volume of space time. The peak regime of where the degrees of freedom maximize out is where the Octonian regime holds. Corresponding to, also , Octonian gravity, when one has flat space, after a significant increase in temperature.**

b) Analog physics, prior to the build up of temperature can be represented by the mappings given by Eq. (53) and Eq. (54) . The first of these mappings is an ergotic mapping , a perfect mixing regime from many universes into our own present universe. By necessity , this mapping requires a deterministic quantum limit as similar to what tHooft included in his embedding of Quantum physics in a larger, non linear theory [33], [34]. This is approximated by current Pilot model build up of an embedding of QM within a more elaborate super structure.

c) The types of discontinuities presented, in Eq. (42), in Eq. (22), Eq. (14), Eq. (15) are ways to[10],[11], [12], [13] investigate the necessity of $\left| \frac{\eta}{s} \approx \varepsilon^{+} \right| << \frac{1}{4\pi}$ giving only $\eta \neq 0, \varepsilon \to \infty$, instead of $\eta \to 0^+$, with **the later case designating when entropy vanishes, which would correspond to no information from prior universes being transferred. I.e. non zero viscosity corresponding to, with almost infinite energy, of when the approach to Octonionic gravity occurs. The other case when viscosity vanishes would be tantamount to when no information is exchanged.**

Understanding the nature of the ergotic mapping in Eq. (53) and Eq. (54) would allow for a rigorous understanding of **the necessity of** $\left|\frac{\eta}{s} \approx \varepsilon^+\right| << \frac{1}{4\pi}$ giving only $\eta \neq 0, \varepsilon \to \infty$, instead of $\eta \to 0^+$, We hope that understanding these issues allows for determining how K essence physics can contribute to emergent structure, and perhaps massive gravitons and avoid symmetry breaking potentials, as used for the Higgs boson, so mentioned in Eq. (29) and Eq. (30) . In doing so, we see first Analog physics in pre Planckian space time, then, briefly the formation of Digital reality, as paramount in the beginning of inflationary cosmology. The genesis of this reality is from an analog physics foundation.

2nd level of conclusion. Determinining if the NEC is valid is essential as establishing a necessary condition for transferal of information from a prior universe to Today's Cosmos;

How to do this? I.e. how to determine if , as an example there is a thermal, flux from a prior universe carrying prior universe information? We will briefly revisit a first principle introduction as to inflaton fluctuations in the beginning which may be part of how to obtain experimental falsifiable criterion

From Weinberg [15] , we can write , from page 192-93, if an inflaton potential $V(\phi) \sim M^{4+\alpha}\phi^{-\alpha}$ then, the inflaton potential has the fluctuation behavior given by

$$\delta\phi \sim t^{\gamma} \tag{71}$$

Then, this assumes

$$\gamma = -.25 \pm \sqrt{\frac{1}{16} - \frac{(6+\alpha)\cdot(1+\alpha)}{(2+\alpha)^2}} \tag{72}$$

The resulting contributions to the CMBR, if worked out, and also connections to gravitational wave astronomy as can be gleaned eventually can be used to pin point an eventual CMBR physics behavior as referred to by Beckwith [42] may after time start giving us ideas if the NEC holds, or does not hold.

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